

IRAS 22198+6336: DISCOVERY OF AN INTERMEDIATE-MASS HOT CORE

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ABSTRACT

We present new Submillimeter Array and Plateau de Bure Interferometer observations of the intermediate-mass object IRAS 22198+6336 in the millimeter continuum and in several molecular line transitions. The millimeter continuum emission reveals a strong and compact source with a mass of $\sim 5 M_{\odot}$ and with properties of Class 0 objects. CO emission shows an outflow with a quadrupolar morphology centered on the position of the dust condensation. The CO outflow emission seems to come from two distinct outflows, one of them associated with SiO outflow emission. A large set of molecular lines has been detected toward a compact dense core clearly coincident with the compact millimeter source, and showing a velocity gradient perpendicular to the outflow traced by CO and SiO. The chemically rich spectrum and the rotational temperatures derived from CH₃CN and CH₃OH (100–150 K) indicate that IRAS 22198+6336 is harboring one of the few intermediate-mass hot cores known at present.

Key words: ISM: individual objects (IRAS 22198+6336) – ISM: lines and bands – radio continuum: ISM – stars: formation

1. INTRODUCTION

Early stages of massive star formation have numerous observational signatures, including molecular masers and outflows, submillimeter continuum sources, hot molecular cores (HMCs), and hyper/ultracompact HII regions. HMCs, first discovered in the vicinity of massive protostars (see Kurtz et al. 2000 for a review), are compact (≤ 0.1 pc, $n \gtrsim 10^7$ cm⁻³) objects with relatively high temperatures ($T_k \gtrsim 100$ K) that show a very rich chemistry in complex organic molecules (CH₃CN, CH₃OH, CH₃OCHO,...).

Recently, regions characterized by high temperatures and densities have also been detected around low-mass protostars (see Ceccarelli 2004 for a review). These so-called hot corinos share many characteristics with HMCs, although the former are $\sim 10^2$ times less massive. Many questions remain unanswered regarding both the high-mass and low-mass hot core phase. For example, HMCs are relatively large objects (~ 0.1 pc) and harbor clusters of stars, while hot corinos seem to be smaller molecular envelopes surrounding a single low-mass young stellar object (YSO) or binary. Additionally, although HMCs are generally understood to be illuminated by the newly born high-mass stars embedded within the cores themselves (Cesaroni 2005), an alternative scenario has been proposed for hot corinos in which jets and shocks are responsible for enhancing the temperature and the abundances of many species (Chandler et al. 2005).

Intermediate-mass YSOs (IMYSOs; $M_{\star} \sim 2$ – $8 M_{\odot}$) are excellent targets to study the different hot core heating mechanisms (radiative or shocks), because they share some properties of massive stars, such as their association with clusters and their ability to photodissociate and ionize the surrounding gas, and present the advantages of being located closer (≤ 1 kpc) than most of the massive star-forming regions and in regions of less complexity. However, there are only two IMYSOs with hot core emission reported in the literature, NGC 7129-FIRS 2 and IC

1396 N (Fuente et al. 2005, 2009), and few conclusions can be inferred from such a small number of intermediate-mass hot cores (IMHCs).

IRAS 22198+6336 (hereafter I22198) is an IMYSO located at a distance of 764 pc (Hirota et al. 2008) with a bolometric luminosity of $370 L_{\odot}$. The presence of a strong and compact submillimeter dust condensation (Jenness et al. 1995) with no near-infrared nor mid-infrared emission suggests that it is a deeply embedded object, classified as an intermediate-mass Class 0 source by Sánchez-Monge et al. (2008). Single-dish telescope observations reveal a CO outflow well centered on the position of the dust condensation (Zhang et al. 2005). Sánchez-Monge et al. (2008) detected slightly resolved partially optically thick centimeter emission consistent with a thermal ionized wind (or radiojet). Additionally, NH₃ and CS dense gas emission together with H₂O and OH maser emission has been detected toward I22198 (Tafalla et al. 1993; Larionov et al. 1999; Valdettaro et al. 2002; Edris et al. 2007; Hirota et al. 2008), all indicative of a very early evolutionary stage.

In this Letter, we present new interferometric continuum and molecular line observations that reveal that I22198 is associated with an IMHC.

2. OBSERVATIONS

We observed the I22198 region with the Submillimeter Array (SMA⁷; Ho et al. 2004) in the 1.3 mm (230 GHz) band using the compact array configuration on 2008 June 11. The phase center was α (J2000) = 22^h21^m27.617 and δ (J2000) = +63°51'42".18, and the projected baselines ranged from 7 kλ to 100 kλ. The two sidebands of the SMA covered the

⁷ The SMA is a joint project between the Smithsonian Astrophysical Observatory and the Academia Sinica Institute of Astronomy and Astrophysics, and is funded by the Smithsonian Institution and the Academia Sinica.

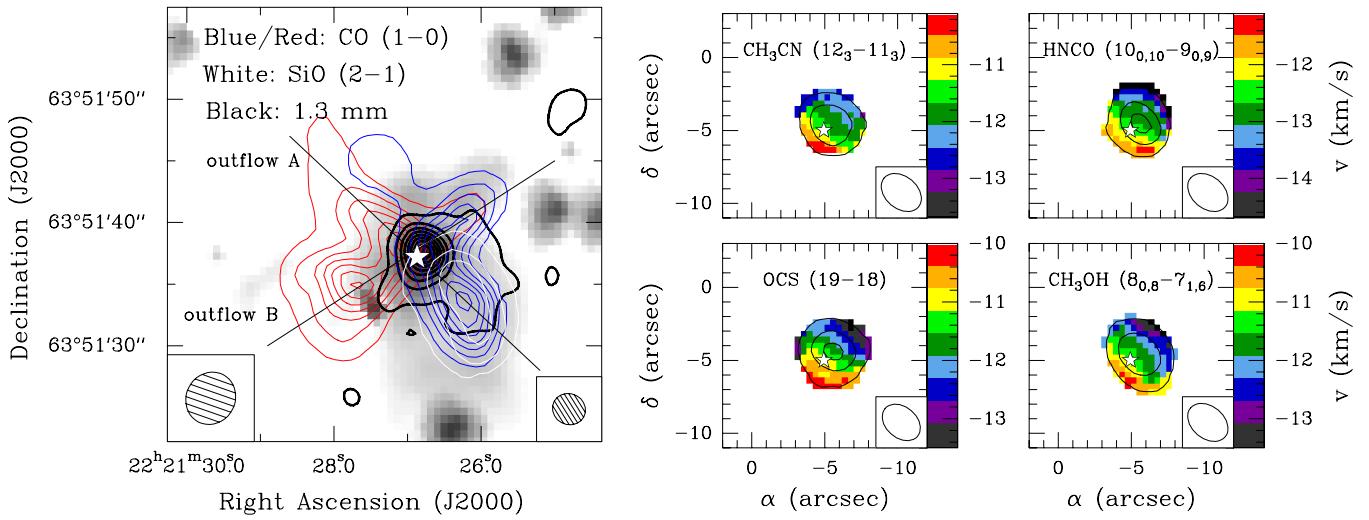


Figure 1. IRAS 22198+6336. Left: blue and red contours: PdBI CO(1–0) intensity map for blueshifted velocities (from -5 to -15 km s^{-1} with respect to the systemic velocity, $v_{\text{LSR}} \simeq -11.3 \text{ km s}^{-1}$) and redshifted velocities (from $+5$ to $+15 \text{ km s}^{-1}$ with respect to the v_{LSR}). Levels start at 5%, increasing in steps of 15% of the peak temperature, 7200 K km s^{-1} and 5400 K km s^{-1} for blue and red, respectively. White contours: PdBI SiO (2–1) intensity map for blueshifted velocities (from -5 to -15 km s^{-1} with respect to the v_{LSR}). Levels start at 5%, increasing in steps of 30% of the peak temperature, 450 K km s^{-1} . Synthesized beams for CO and SiO images are shown in the bottom right and left corners, respectively. Black contours: SMA 1.3 mm continuum map. Levels are -4 , and 4 to 60 in steps of 8, times $3.92 \text{ mJy beam}^{-1}$. Gray scale: $4.5 \mu\text{m}$ IRAC/*Spitzer* image. Right: each of the four panels show in black contours the zero-order moment (intensity) map and in color scale the first-order moment (velocity) map for four molecular transitions observed with the SMA. Levels are 5%, 45%, and 85% with the intensity peaks $2.67, 3.20, 3.32, 3.97 \text{ km s}^{-1}$, for CH_3CN , HNCO , OCS , and CH_3OH , respectively. The $(0, 0)$ position corresponds to the SMA phase center. The white star marks the position of the centimeter continuum source detected by Sánchez-Monge et al. (2008).

frequency ranges of 219.4 – 221.4 GHz and 229.4 – 231.4 GHz , with a spectral resolution of $\sim 0.5 \text{ km s}^{-1}$. System temperatures ranged between 200 and 300 K . The zenith opacities at 225 GHz were around 0.15 and 0.20 during the track. The FWHM of the primary beam at 1.3 mm was $\sim 56''$. Bandpass calibration was performed by observing the quasar $3\text{C}454.3$. Amplitude and phase calibrations were achieved by monitoring $0019+734$ and $1928+739$, resulting in an rms phase of $\sim 30^\circ$. The absolute flux density scale was determined from Callisto and Jupiter with an estimated uncertainty around 15%. Data were calibrated and imaged with the MIRIAD software package. The continuum was constructed in the (u, v) domain from the line-free channels. Imaging was performed using natural weighting, obtaining a synthesized beam of $3''.0 \times 2''.2$ with a P.A. = 50° , and 1σ rms of 80 mJy beam^{-1} per channel, and $3.9 \text{ mJy beam}^{-1}$ for the continuum.

The IRAM Plateau de Bure Interferometer (PdBI)⁸ was used to observe the CO(1–0) (115.12 GHz) and SiO (2–1) (86.85 GHz) molecular transitions toward I22198. We carried out two-pointing mosaic observations (primary beam $\simeq 56''$), in 2008 and 2009, with the array in the C and D configurations. The two spectral setups, tuned at 115.27 GHz (2.7 mm) and at 86.85 GHz (3.5 mm), include several molecular transitions, which will be presented in a forthcoming paper (Á. Sánchez-Monge et al. 2010, in preparation). For the CO and SiO, we used a correlator unit of 40 MHz of bandwidth with 512 spectral channels, which provides a spectral resolution of $\sim 0.3 \text{ km s}^{-1}$. A number of 320 MHz continuum units were used to image the continuum, six at 115 GHz and three at 86 GHz . The typical system temperatures for the receivers were $\sim 200 \text{ K}$ at 115 GHz and $\sim 100 \text{ K}$ at 86 GHz . Bandpass calibration was performed by observing $3\text{C}273$ and $3\text{C}454.3$. Amplitude

and phase calibrations were achieved by monitoring $0116+731$, $1928+738$, and $2037+511$, resulting in a phase rms of around 25° . The absolute flux density scale was determined from $\text{MWC}349$ and $3\text{C}84$, with an estimated uncertainty around 25%. The data were calibrated with the program CLIC and imaged with MAPPING, both parts of the GILDAS⁹ software package. Imaging was performed using natural weighting, obtaining a synthesized beam of $3''.0 \times 2''.5$ with a P.A. = 107° at 115 GHz and $4''.5 \times 4''.0$ with a P.A. = 81° at 87 GHz . The 1σ rms for the channel and continuum images were 40 and $0.6 \text{ mJy beam}^{-1}$ at 115 GHz , and 30 and $0.3 \text{ mJy beam}^{-1}$ at 87 GHz , respectively.

3. RESULTS

3.1. Continuum Results

In Figure 1 (left panel), we show the interferometric millimeter continuum map toward I22198 (black contours). We detected a compact source (with coordinates $\alpha = 22^{\text{h}}21^{\text{m}}26.^{\text{s}}78$ and $\delta = +63^\circ 51' 37.5''$) coincident with the centimeter source reported by Sánchez-Monge et al. (2008) and surrounded by a faint structure extended toward the southwest. The source is clearly detected at $3.5, 2.7$, and 1.3 mm with integrated flux densities of $24 \pm 6 \text{ mJy}$, $55 \pm 15 \text{ mJy}$, and $500 \pm 75 \text{ mJy}$, respectively, and has a similar morphology at the three wavelengths. The deconvolved size of the source is $3''.0 \times 2''.1$ with a P.A. = 150° , corresponding to $2300 \times 1600 \text{ AU}$. The millimeter continuum emission is likely tracing the dust envelope of I22198.

The spectral energy distribution (SED) reported by Sánchez-Monge et al. (2008) has been improved with the new continuum data. We also included the fluxes between $3.6 \mu\text{m}$ and $8.0 \mu\text{m}$ extracted from the *Spitzer* Space Telescope Infrared Array Camera (IRAC) maps (PI: Giovanni Fazio, project:

⁸ The Plateau de Bure Interferometer (PdBI) is operated by the Institut de Radioastronomie Millimétrique (IRAM), which is supported by INSU/CNRS (France), MPG (Germany), and IGN (Spain).

⁹ GILDAS: Grenoble Image and Line Data Analysis System; see <http://www.iram.fr/IRAMFR/GILDAS>.

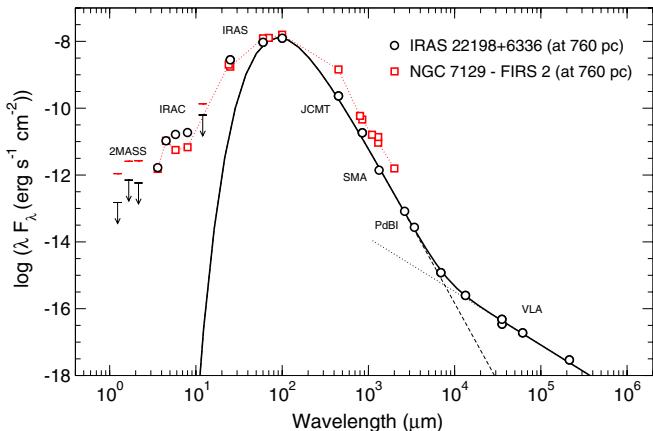


Figure 2. Spectral energy distribution of IRAS 22198+6336 (black circles). Solid line: sum of the modified blackbody law and the thermal ionized wind. Black dashed line: modified blackbody law for the dust envelope with a dust emissivity index of $\beta = 1.7$, a source radius of $3'0$, a dust temperature of 36 K, a dust mass of $5 M_{\odot}$, and a dust mass opacity coefficient of $0.9 \text{ cm}^2 \text{ g}^{-1}$ at 1.3 mm. Black dotted line: thermal free-free emission from an ionized wind with a spectral index of ~ 0.6 at centimeter wavelengths. We estimated the total bolometric luminosity, from IR up to mm wavelengths, to be $L_{\text{bol}} \simeq 370 L_{\odot}$ for a distance of 0.76 kpc. Red squares: NGC 7129-FIRS 2 data (from Eiroa et al. 1998, Crimier et al. 2010, and 2MASS and IRAC catalogs) rescaled at the distance of I22198 (760 pc). We estimated $L_{\text{bol}} \simeq 480 L_{\odot}$. Labels of telescopes refer only to I22198 data.

40147). We retrieved the images using the Leopard¹⁰ software and evaluated the fluxes using the MOPEX¹¹ software. The global SED (open circles in Figure 2) was fitted by a modified blackbody law plus a thermal ionized wind (with a spectral index of ~ 0.6 at centimeter wavelengths). Assuming a dust mass opacity coefficient given by $\kappa_v = \kappa_0(v/v_0)^{\beta}$ with $\kappa_0 = 0.9 \text{ cm}^2 \text{ g}^{-1}$ at $v_0 = 230 \text{ GHz}$ (Ossenkopf & Henning 1994), the SED is well fitted with a dust emissivity index of 1.7 ± 0.2 , a dust temperature of $36 \pm 2 \text{ K}$, and an envelope mass of $5 \pm 1 M_{\odot}$. As shown in Figure 2, the SED of I22198 resembles that of Class 0 objects (André et al. 1993).

3.2. Molecular Results

Interferometric $^{12}\text{CO}(1-0)$ (Figure 1 (left)) and $^{12}\text{CO}(2-1)$ (Á. Sánchez-Monge et al. 2010, in preparation) maps reveal an outflow with a quadrupolar morphology clearly centered on the position of the dust condensation. The outflow emission spans a velocity range from -25 km s^{-1} to $+5 \text{ km s}^{-1}$ for the $J = 1 \rightarrow 0$ transition (the systemic velocity is $v_{\text{LSR}} \simeq -11.3 \text{ km s}^{-1}$). The quadrupolar morphology of the outflow can be interpreted as the walls of a single outflow or as the superposition of two bipolar outflows: outflow A in the southwest–northeast direction and outflow B in the northwest–southeast direction. Both outflows are centered on the position of the millimeter continuum source, which could be a millimeter clump harboring at least two sources. In fact, higher angular resolution observations ($0''.4$ or 300 AU) with the PdBI at 1.3 mm reveal a strong source with indications of a faint extension to the south (A. Palau et al. 2010, in preparation). In the following, we will assume the two outflow interpretation. We estimated the physical parameters of outflows A and B (see Table 1) and found values similar to the outflow parameters found in other IMYSOs (e. g., IC 1396 N: Beltrán et al. 2002; IRAS 22272+6358 A: Beltrán et al. 2006).

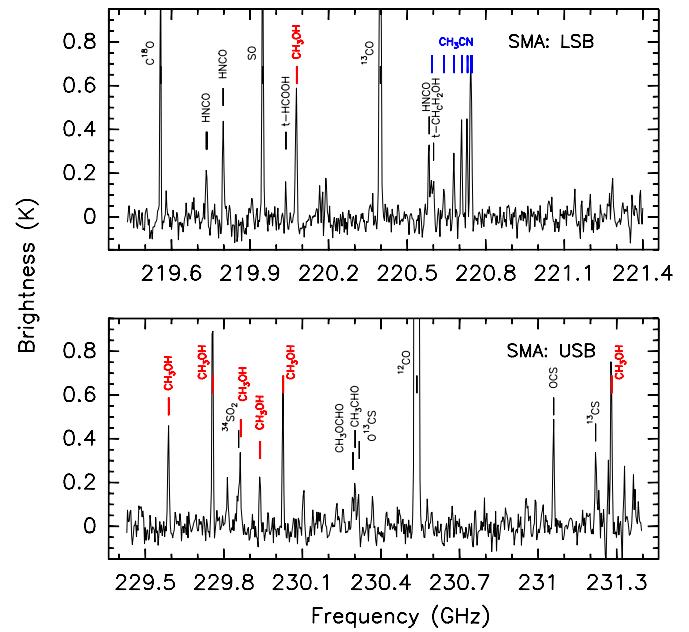


Figure 3. 230 GHz continuum-free spectra in the image domain from the SMA data, for the lower side band (LSB) in the top panel and the upper side band (USB) in the bottom panel. The molecular transitions used in the rotational diagrams (Figure 4) are shown in blue and red.

The integrated SiO(2–1) emission, spanning a velocity range from -25 to -8 km s^{-1} , is shown in white contours in Figure 1. This emission coincides with the blueshifted lobe of outflow A. No redshifted counterpart is observed for the SiO emission, however. The detection of SiO in outflow A suggests that it is younger than outflow B, because gas-phase SiO abundances may decrease with time (Klaassen & Wilson 2007; Shang et al. 2006). The SiO detection also suggests that the CO quadrupolar morphology arises from two distinct outflows, because SiO generally traces material associated with the primary jet driving the outflow rather than the walls of the outflow cavity (e. g., H211: Palau et al. 2006; Lee et al. 2007). We also note that the IRAC 4.5 μm band (gray scale in Figure 1 (left)) shows extended emission (“green fuzzy”) associated mainly with the blueshifted CO outflow lobes. Such 4.5 μm extended emission is thought to be associated with shocked H₂ emission (De Buizer & Vacca 2010).

The wide-band SMA spectrum toward I22198 (Figure 3) reveals a chemically rich dense core with typical hot core features (e. g., Cesaroni et al. 1999; Gibb et al. 2000). Molecular line transitions detected with $T_B \gtrsim 0.2 \text{ K}$ have been identified and imaged. We note that the line widths observed for most of the molecules are $5\text{--}8 \text{ km s}^{-1}$. Figure 1 (right panel) shows the maps of high-density tracers with a compact source ($2''.1 \times 1''.3$ with P.A. = 15°) at the position of the millimeter continuum source. Although the maps of each molecule show a source barely resolved, and higher angular resolutions should be made, first-order moment (velocity field) maps for these high-density tracers (Figure 1 (right)) show velocity gradients in the northwest–southeast direction, perpendicular to outflow A, presumably the youngest outflow. Most of these molecules trace disks and/or toroids in other massive YSOs (e. g., Beltrán et al. 2006). However, the presence of outflow B perpendicular to outflow A makes the interpretation of these velocity gradients uncertain. Assuming that these gradients trace rotation, the

¹⁰ <http://ssc.spitzer.caltech.edu/warmmission/propkit/spot/>

¹¹ <http://ssc.spitzer.caltech.edu/dataanalysistools/tools/mopex/>

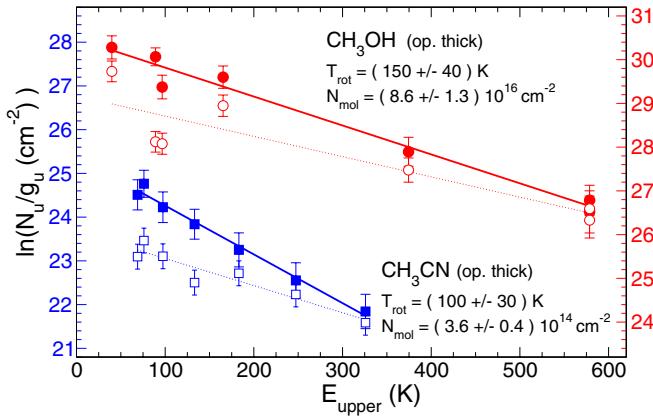


Figure 4. CH_3CN (blue squares) and CH_3OH (red circles) rotational diagrams. Open symbols and dashed lines refer to the optically thin approximation, while filled symbols and solid lines refer to the opacity which yields the best linear fit. For the optically thick fit we give the values in the figure and adopt them in the main text. For the optically thin fit, $T_{\text{rot}} = 160 \pm 30$ K and $N_{\text{mol}} = (1.7 \pm 0.3) \times 10^{14} \text{ cm}^{-2}$ for CH_3CN , and $T_{\text{rot}} = 220 \pm 60$ K and $N_{\text{mol}} = (4.6 \pm 1.6) \times 10^{16} \text{ cm}^{-2}$ for CH_3OH . Vertical left and right axes refer to CH_3CN and CH_3OH , respectively.

dynamical mass has been calculated from the expression $M_{\text{dyn}} = v_{\text{rot}}^2 R_{\text{rot}} / G \sin^2 i$, where v_{rot} is the velocity estimated from the gradient, R_{rot} is the radius of the core, G is the gravitational constant, and i is the inclination angle assumed to be 90° (edge-on). In our case, with $v_{\text{rot}} \simeq 2 \text{ km s}^{-1}$ and $R_{\text{rot}} \simeq 1''$ for the four molecules shown in Figure 1 (right), the dynamical mass is $\sim 3.5 M_\odot$. Additionally, following the models of Palla & Stahler (1993) and assuming a luminosity for the object of $\sim 370 L_\odot$ (Figure 2), the mass of the embedded YSO is about $4-5 M_\odot$. Thus, the mass of the internal YSO estimated from both methods is $\sim 4 M_\odot$, and this is similar to the dust envelope mass ($\sim 5 M_\odot$), which is characteristic of Class 0 objects.

Because of the wide SMA band, one can observe several CH_3CN and CH_3OH transitions simultaneously. Following the rotational diagram method (Goldsmith & Langer 1999; Araya et al. 2005), which assumes that all molecular levels are populated according to the same excitation temperature, we can derive this temperature, T_{rot} , and the total column density, N_{mol} , for these molecules. We constructed the rotational diagrams following two methods. First, we assumed optically thin emission (open symbols in Figure 4). Second, we estimated the opacity for each transition (following Goicoechea et al. 2006; see also Girart et al. 2002) and adopted the value that yielded the best linear fit (filled symbols in Figure 4). The resulting opacities for the different CH_3CN and CH_3OH transitions range from 0.3 to 3.5 and from 0.2 to 6.5, respectively. In the discussion, we will use the results of the latter fit, which are $T_{\text{rot}} = 100 \pm 30$ K and $N_{\text{mol}} = (3.6 \pm 0.4) \times 10^{14} \text{ cm}^{-2}$ for CH_3CN , and $T_{\text{rot}} = 150 \pm 40$ K and $N_{\text{mol}} = (8.6 \pm 1.3) \times 10^{16} \text{ cm}^{-2}$ for CH_3OH . The high temperatures, together with the chemically rich spectrum, are clear evidence that I22198 is an IMHC.

4. DISCUSSION AND CONCLUSIONS

Our interferometric continuum and molecular line observations toward I22198 reveal a compact and strong millimeter source driving a quadrupolar outflow, which is associated with a compact dense core with hot core signatures. At infrared wavelengths, IRAC images reveal a single source coincident with the hot core and with “green fuzzies” at $4.5 \mu\text{m}$. The high rotational temperature ($100-150$ K) estimated from the CH_3CN and

Table 1
Properties of Intermediate-mass Hot Cores^a

Property	IRAS 22198+6336	NGC 7129-FIRS 2	IC 1396 N
Distance (pc)	760	1250	750
$L_{\text{bol}} (L_\odot)$	370	480	<300
$M_{\text{env}} (M_\odot)$	5	2	<5
Clustering (mm)	No	No	Yes
IRAC source	Yes	Yes	No
$L_{\text{cm}} (\text{mJy kpc}^2)$	0.34	<0.36	<0.15
Outflow	Yes	Yes	Yes
Hot core information			
T_{rot} (K)	100–150	50–100	100
N_{H_2} (cm^{-2})	4×10^{23} ^b	8×10^{24}	3×10^{24}
$X_{\text{CH}_3\text{CN}}$	1×10^{-9}	7×10^{-9}	5×10^{-10}
Outflow information			
t_{dyn} (yr)	1300/1200	4000	2600
\dot{M} ($10^{-6} M_\odot \text{ yr}^{-1}$)	9/8	7	50
P ($M_\odot \text{ km s}^{-1}$)	0.3/0.2	0.4	4

Notes.

^a References are: I22198: Sánchez-Monge et al. 2008; this work (the two values in outflow parameters correspond to outflows A and B, respectively); NGC 7129-FIRS 2: Fuente et al. 2001, 2005; Crimier et al. 2010 (centimeter data from the VLA project AR304, and outflow parameters: Á. Sánchez-Monge 2010, unpublished); IC 1396 N: Beltrán et al. 2002; Neri et al. 2007; Fuente et al. 2009. Upper limits for IC 1396 N are due to the cluster properties (see the text).

^b Estimated following Fuente et al. (2005), assuming $T = 100$ K and a size of $2''.1 \times 1''.3$, which is the deconvolved size of the CH_3CN core.

CH_3OH rotational diagrams, and the chemically rich spectrum, make I22198 one of the few IMHCs known at present. There are only two other IMHCs reported in the literature: NGC 7129-FIRS 2 (Fuente et al. 2005) and IC 1396 N (Neri et al. 2007; Fuente et al. 2009). Table 1 lists the main properties of the three IMHCs known to date.

IC 1396 N is associated with a single *IRAS* source, IRAS 21319+5802. However, high angular resolution millimeter images reveal a cluster of sources, only one of which is associated with the IMHC (Neri et al. 2007). Moreover, while *Spitzer*/IRAC observations reveal a counterpart for the millimeter source 41.73+12.8, there is no counterpart for the IMHC millimeter source (41.86+11.9; Choudhury et al. 2010; Neri et al. 2007). All this makes it difficult to assess the IMHC contribution to the *IRAS* and IRAC fluxes, and hence we refrain from comparing IC 1396 N with I22198.

The other IMHC, NGC 7129-FIRS 2, has a bolometric luminosity similar to I22198, has a single associated source dominating the millimeter emission (Fuente et al. 2005), and also is clearly associated with a source detected in IRAC and MIPS bands (Á. Sánchez-Monge 2010, unpublished). Eiroa et al. (1998) study the far-infrared emission and find one HIRES source, which is coincident with the MIPS source (Crimier et al. 2010). The association of the millimeter, far-infrared, and mid-infrared sources with the IMHC is clear, as in the case of I22198 and, similarly, “green fuzzies” in the $4.5 \mu\text{m}$ IRAC image appear associated with the molecular outflow, which is also quadrupolar (Fuente et al. 2001) and has physical parameters similar to the outflows in I22198 (Table 1). Regarding the centimeter emission, the upper limit for NGC 7129-FIRS 2 is consistent with the centimeter emission found toward I22198 (Table 1). Furthermore, Figure 2 shows that both IMHCs have quite similar Class 0-like SEDs. Thus, NGC 7129-FIRS 2 and I22198 have very similar properties

concerning clustering, centimeter, and outflow emission, and have a similar SED. Moreover, the hot core temperature, T_{rot} , and CH₃CN abundance, $X_{\text{CH}_3\text{CN}}$, are similar for these two IMHCs. At this point, our data do not allow us to distinguish between the proposed hot core scenarios (radiative or shocked), because the two IMHCs have similar luminosities and outflow/centimeter properties. A detailed study of different molecular abundances (enhanced by the radiation field or by shocks) will be presented in a forthcoming paper and should clarify the formation scenario of the hot core.

In summary, the properties of the newly discovered IMHC, I22198, and its comparison with the other known IMHCs help us better establish the characteristics of the hot core phase for IMYSOs. We have found that the IMHC phase is coeval with mass ejection (as shown by outflow, “green fuzzies,” and faint centimeter emission), presents Class 0-like SEDs, has hot core temperatures ~ 100 K, and $X_{\text{CH}_3\text{CN}} \simeq 10^{-9}$. A detailed comparison of I22198 and NGC 7129-FIRS 2 and the discovery of new IMHCs will provide fundamental tools to understand the hot core phase of star formation.

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REFERENCES

- André, P., Ward-Thompson, D., & Barsony, M. 1993, *ApJ*, **406**, 122
 Araya, E., Hofner, P., Kurtz, S., Bronfman, L., & DeDeo, S. 2005, *ApJS*, **157**, 279
 Beltrán, M. T., Girart, J. M., & Estalella, R. 2006, *A&A*, **457**, 865
- Ceccarelli, C. 2004, in ASP Conf. Ser. 323, Star Formation in the Interstellar Medium: In Honor of David Hollenbach, ed. D. Johnstone et al. (San Francisco, CA: ASP), 195
- Cesaroni, R. 2005, in IAU Symp. Ser. 227, Massive Star Birth: A Crossroads of Astrophysics, ed. R. Cesaroni et al. (Cambridge: Cambridge Univ. Press), 59
- Cesaroni, R., Felli, M., Jenness, T., Neri, R., Olmi, L., Roberto, M., Testi, L., & Walmsley, C. M. 1999, *A&A*, **345**, 949
- Chandler, C. J., Brogan, C. L., Shirley, Y. L., & Loinard, L. 2005, *ApJ*, **632**, 371
- Choudhury, R., Mookerjea, B., & Bhatt, H. C. 2010, *ApJ*, **717**, 1067
- Crimier, N., et al. 2010, *A&A*, **516**, A102
- De Buizer, J. M., & Vacca, W. D. 2010, *AJ*, **140**, 196
- Edris, K. A., Fuller, G. A., & Cohen, R. J. 2007, *A&A*, **465**, 865
- Eiroa, C., Palacios, J., & Casali, M. M. 1998, *A&A*, **335**, 243
- Fuente, A., Neri, R., & Caselli, P. 2005, *A&A*, **444**, 481
- Fuente, A., Neri, R., Martín-Pintado, J., Bachiller, R., Rodríguez-Franco, A., & Palla, F. 2001, *A&A*, **366**, 873
- Fuente, A., et al. 2009, *A&A*, **507**, 1475
- Gibb, E., Nummelin, A., Irvine, W. M., Whittet, D. C. B., & Bergman, P. 2000, *ApJ*, **545**, 309
- Girart, J. M., Viti, S., Williams, D. A., Estalella, R., & Ho, P. T. P. 2002, *A&A*, **388**, 1004
- Goicoechea, J. R., Pety, J., Gerin, M., Teyssier, D., Roueff, E., Hily-Blant, P., & Baek, S. 2006, *A&A*, **456**, 565
- Goldsmith, P. F., & Langer, W. D. 1999, *ApJ*, **517**, 209
- Hirota, T., et al. 2008, *PASJ*, **60**, 961
- Ho, P. T. P., Moran, J. M., & Lo, K. Y. 2004, *ApJ*, **616**, L1
- Jenness, T., Scott, P. F., & Padman, R. 1995, *MNRAS*, **276**, 1024
- Klaassen, P. D., & Wilson, C. D. 2007, *ApJ*, **663**, 1092
- Kurtz, S., Cesaroni, R., Churchwell, E., Hofner, P., & Walmsley, C. M. 2000, in *Protostars and Planets IV*, ed. V. Mannings, A. P. Boss, & S. S. Russell (Tucson, AZ: Univ. Arizona Press), 299
- Larionov, G. M., Val'tts, I. E., Winnberg, A., Johansson, L. E. B., Booth, R. S., & Golubev, V. V. 1999, *A&AS*, **139**, 257
- Lee, C.-F., Ho, P. T. P., Palau, A., Hirano, N., Bourke, T. L., Shang, H., & Zhang, Q. 2007, *ApJ*, **670**, 1188
- Neri, R., et al. 2007, *A&A*, **468**, L33
- Ossenkopf, V., & Henning, T. 1994, *A&A*, **291**, 943
- Palau, A., et al. 2006, *ApJ*, **636**, L137
- Palla, F., & Stahler, S. W. 1993, *ApJ*, **418**, 414
- Sánchez-Monge, Á., Palau, A., Estalella, R., Beltrán, M. T., & Girart, J. M. 2008, *A&A*, **485**, 497
- Shang, H., Allen, A., Li, Z.-Y., Liu, C.-F., Chou, M.-Y., & Anderson, J. 2006, *ApJ*, **649**, 845
- Tafalla, M., Bachiller, R., & Martin-Pintado, J. 1993, *ApJ*, **403**, 175
- Valdettaro, R., Palla, F., Brand, J., Cesaroni, R., Comoretto, G., Felli, M., & Palagi, F. 2002, *A&A*, **383**, 244
- Zhang, Q., Hunter, T. R., Brand, J., Sridharan, T. K., Cesaroni, R., Molinari, S., Wang, J., & Kramer, M. 2005, *ApJ*, **625**, 864